

MULTI-WAVELENGTH STUDY OF THE Be/X-RAY BINARY MXB 0656–072

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ABSTRACT

We present and analyze the optical photometric and spectroscopic data of the Be/X-ray binary MXB 0656–072 from 2006 to 2009. A 101.2 day orbital period is found, for the first time, from the present public X-ray data (*Swift*/BAT and *RXTE*/ASM). The anti-correlation between the H α emission and the *UBV* brightness of MXB 0656–072 during our 2007 observations indicates that a mass ejection event took place in the system. After the mass ejection, a low-density region might develop around the Oe star. With the outward motion of the circumstellar disk, the outer part of the disk interacted with the neutron star around its periastron passage and a series of X-ray outbursts were triggered between MJD 54350 and MJD 54850. The Proportional Counter Array-HEXTE spectra during the 2007–2008 X-ray outbursts could be well fitted by a cutoff power law with low-energy absorption, together with an iron line around 6.4 keV, and a broad cyclotron resonance feature around 30 keV. The same variability of the soft and hard X-ray colors in 2.3–21 keV indicated that there were no overall changes in the spectral shape during the X-ray outbursts, which might only be connected with the changes of the mass accretion rate onto the neutron star.

Key words: stars: emission-line, Be – stars: individual (MXB 0656–072) – stars: neutron – X-rays: binaries

Online-only material: color figure

1. INTRODUCTION

Be/X-ray binaries represent the largest subclass of high-mass X-ray binaries, in which a neutron star orbits around a massive and early-type star in a wide and eccentric orbit. Currently, no Be/black hole system has been found (Zhang et al. 2004). There are two different kinds of disks in Be/X-ray binaries, a circumstellar disk around the Be star, and an accretion disk around the neutron star. The formation of a Be disk may be connected with the rapid rotation of the Be star (Porter & Rivinius 2003). Be/X-ray binaries show two different X-ray outburst activities: Type I X-ray outbursts, which are usually associated with the periastron passage of the neutron star through the environment of the circumstellar disk, and Type II X-ray outbursts, which could occur at any orbital phase and last for a large fraction of the orbital period (see Reig 2011 and references therein).

The transient X-ray source MXB 0656–072 was discovered on 1975 September 20 with SAS-3 (Clark et al. 1975). The following high-sensitivity-mode observations showed that the 3–6 keV flux of the source was 0.05 and 0.07 times that of the Crab Nebula on 1976 March 19 and March 27, respectively (Kaluzienski et al. 1976). According to its long-term variability in X-rays, MXB 0656–072 was classified as a low-mass X-ray binary by Liu et al. (2001).

RXTE/ASM detected a new X-ray outburst of MXB 0656–072 in 2003 October (Remillard & Marshall 2003) that reached a peak X-ray luminosity of 200 mCrab. *RXTE*/PCA observations on 2003 October 19–20 indicated that the source was a pulsar with a spin period of 160.7 s (Morgan et al. 2003). Using *ROSAT* PSPC observations, Pakull et al. (2003) identified the optical counterpart and revealed an O9.7 Ve spectral type. Thus, MXB 0656–072 was classified as a Be/X-ray binary in a recent high-mass X-ray binary catalog (Liu et al. 2006).

McBride et al. (2006) analyzed the *RXTE* observations during the 2003 outburst and found a cyclotron resonance scattering feature at 32.8 keV and an average pulse period of 160.4 ± 0.4 s with a spin-up of 0.45 s across the outburst. They also estimated a distance of 3.9 ± 0.1 kpc to the source from the observed *V* magnitude ($V = 12.05\text{--}12.38$) and *B* – *V* color ($B - V = 1.02\text{--}0.86$) (Pakull et al. 2003).

A series of X-ray outbursts that occurred between 2007 November 10 and 2008 November 10 was reported by *International Gamma-Ray Astrophysics Laboratory* (Kreykenbohm et al. 2007), *RXTE*/ASM (Pottschmidt et al. 2007), and *Swift*/BAT (Kennea et al. 2007). Yan et al. (2007) reported their spectroscopic observations on MXB 0656–072 and found a strong H α emission line before its 2007 X-ray outburst.

In this paper, we present our optical photometric and spectroscopic observations on MXB 0656–072 from 2005 to 2009 and simultaneous X-ray observations in Section 2. The optical and X-ray variability during different X-ray states is analyzed in Section 3. In Section 4, we discuss the X-ray periodic outburst nature of the system. The results are summarized in Section 5.

2. OBSERVATIONS

The multi-wavelength observations, including the intermediate-resolution optical spectroscopy, *UBV* photometry, and X-rays, are described in the following subsections.

2.1. Optical Spectroscopy

We obtained several spectra of MXB 0656–072 with the 2.16 m telescope at Xinglong Station of National Astronomical Observatories, China (NAOC), from 2005 to 2009. The optical spectroscopy with an intermediate resolution of $1.22 \text{ \AA pixel}^{-1}$ was performed with a CCD grating spectrograph at the Cassegrain focus of the telescope. We took blue

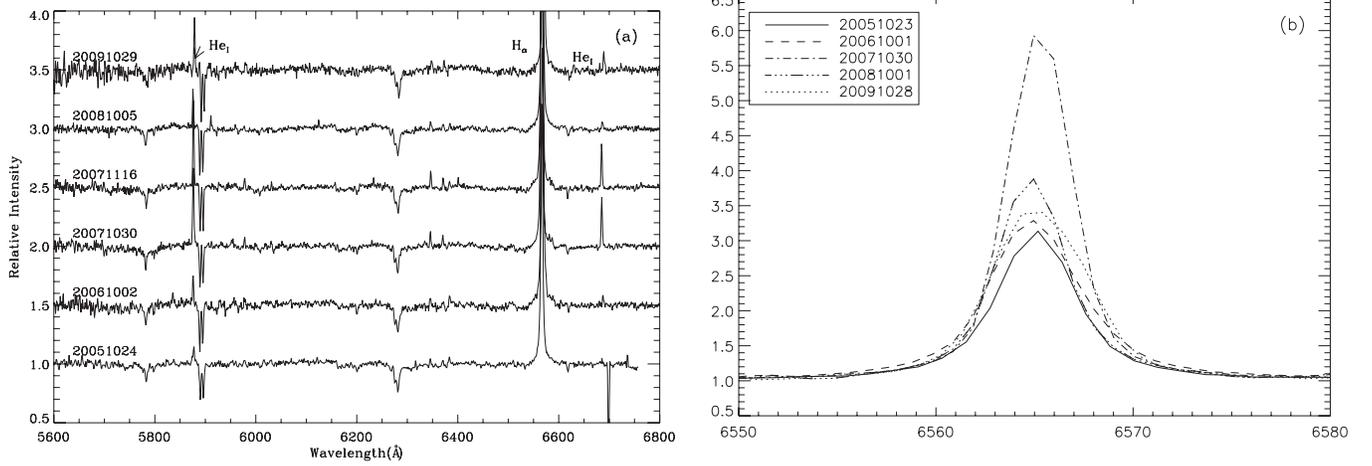


Figure 1. (a) Selected spectra from our 2005 to 2009 observations. All spectra have the continuum level normalized and offset vertically to allow direct comparison. The observational date is marked on the left part of each spectrum in the YYYYMMDD format. The dip on the right part of spectrum on 2005 October 24 was caused by bad pixels on the CCD. (b) The zoomed panel of the left figure in the $H\alpha$ region with all spectra at the same continuum level.

Table 1
Summary of the Spectroscopic Observations of MXB 0656–072

Date (YYYYMMDD)	MJD	Exposure (s)	EW (\AA)	Error (\AA)
20051023	53666.88	1000.0	14.44	0.13
20051023	53666.89	1000.0	14.20	0.38
20051024	53667.89	1000.0	14.05	0.16
20060928	54006.85	1200.0	16.81	0.13
20060929	54007.85	1200.0	16.02	0.31
20061001	54009.88	1200.0	17.03	0.21
20061002	54010.88	900.0	17.14	0.32
20071028	54401.91	120.0	24.34	0.45
20071030	54403.88	1800.0	25.06	0.30
20071031	54404.85	1800.0	25.25	0.16
20071101	54405.87	1800.0	25.41	0.34
20071116	54420.77	1800.0	23.75	0.23
20071116	54420.84	2400.0	24.16	0.22
20080930	54739.85	1800.0	16.88	0.36
20081001	54740.84	1800.0	16.38	0.31
20081005	54744.86	1800.0	16.19	0.09
20081006	54745.85	1800.0	16.35	0.18
20081009	54748.86	1800.0	16.38	0.09
20091023	55127.90	1200.0	16.92	0.34
20091024	55128.87	1500.0	16.96	0.14
20091025	55129.86	1800.0	15.97	0.16
20091028	55132.90	1800.0	17.91	0.12
20091029	55133.86	2100.0	17.73	0.11

and red spectra covering the wavelength ranges 4300–5500 and 5500–6700 \AA , respectively, at different times. All spectra were reduced with the IRAF⁴ package. The data were bias-subtracted and flat-field-corrected, and the cosmic rays were removed. Helium–argon spectra were taken in order to obtain the pixel–wavelength relation. To improve this relation, we also used the diffuse interstellar bands at 6614 and 6379 \AA observed in the spectra.

Some selected and normalized red spectra from 2005 to 2009 and a zoom of the $H\alpha$ region are shown in Figures 1(a) and (b), respectively. The corresponding observational dates are marked in the left part of each spectrum. The equivalent

Table 2
Log of the UBV Optical Photometric Observations of MXB 0656–072

Date	Exposures [$U/B/V$]	Seeing (arcsec)
2007 Oct 30	480s+360s/240s+120s/2 \times 120s	2.0
2007 Oct 31	360s/120s/60s	2.0
2007 Nov 1	2 \times 360s/2 \times 120s/2 \times 60s	3.5
2008 Oct 5	3 \times 480s/3 \times 120s/3 \times 60s	3.5
2008 Oct 6	3 \times 480s/3 \times 120s/3 \times 60s	3.5
2008 Oct 8	4 \times 480s/4 \times 120s/2 \times 60s	2.0
2009 Oct 27	4 \times 540s/4 \times 120s/4 \times 60s	4.7
2009 Oct 28	2 \times 540s/2 \times 120s/2 \times 60s	2.3

width (EW) of the $H\alpha$ line has been measured by selecting a continuum point on each side of the line and integrating the flux relative to the straight line between the two points using the procedures available in IRAF. The measurements were repeated five times for each spectrum and the error was estimated from the distribution of obtained values. The EW($H\alpha$) typical error is within 3%. This error arises due to the subjective selection of the continuum. The EW($H\alpha$) results are listed in Table 1 together with the observational date, the MJD, and the exposure.

The blue spectra were taken as well during our five years of observations (see Figure 2). He I λ 5016, $H\beta$, and $H\gamma$ are clearly seen. Both observational date and EW($H\beta$) are marked at the left and right part of each spectrum, respectively.

2.2. UBV Photometry

Since 2007, we performed systematic photometric observations on a sample of X-ray binaries with both the 100 cm Education and Science Telescope (EST) and the 80 cm Tsinghua-NAOC Telescope (TNT) at Xinglong Station of NAOC. The EST, manufactured by EOS Technologies, is an altazimuth-mounted reflector with Nasmyth foci at a focal ratio of $f/8$. TNT is an equatorial-mounted Cassegrain system with a focal of $f/10$, made by AstroOptik, funded by Tsinghua University in 2002, and jointly operated with NAOC. Both telescopes are equipped with the same type of Princeton Instrument 1340×1300 thin back-illuminated CCD. The CCD cameras use standard Johnson-Cousins $UBVRI$ filters made by Custom Scientific. Table 2 reports the log of our photometric observations

⁴ IRAF is distributed by NOAO, which is operated by the Association of Universities for Research in Astronomy, Inc., under cooperation with the National Science Foundation.

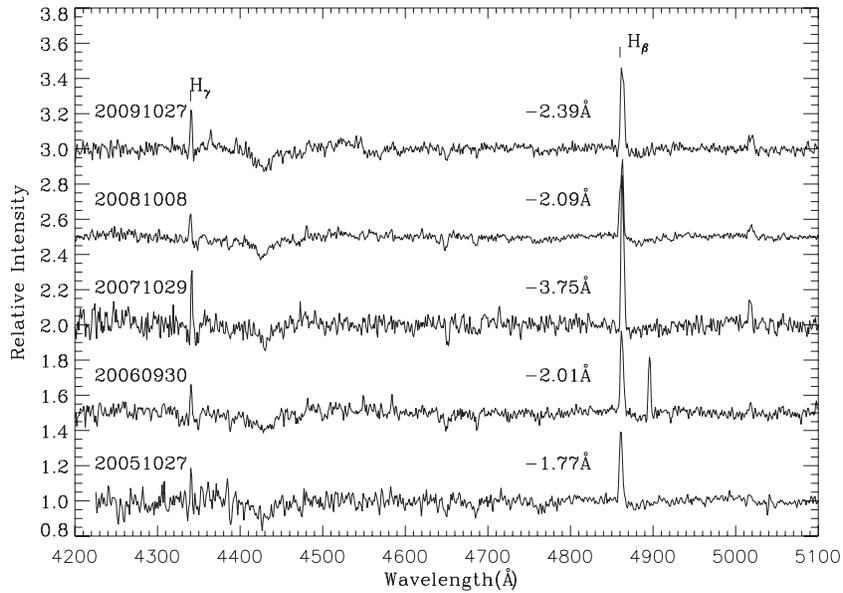


Figure 2. Blue spectra of MXB 0656–072 from 2005 to 2009. All spectra have the continuum level normalized and offset vertically to allow direct comparison. Every observational date and $EW(H\beta)$ are reported for each spectrum. The spike near the $H\beta$ line on 2006 September 30 is caused by a cosmic ray.

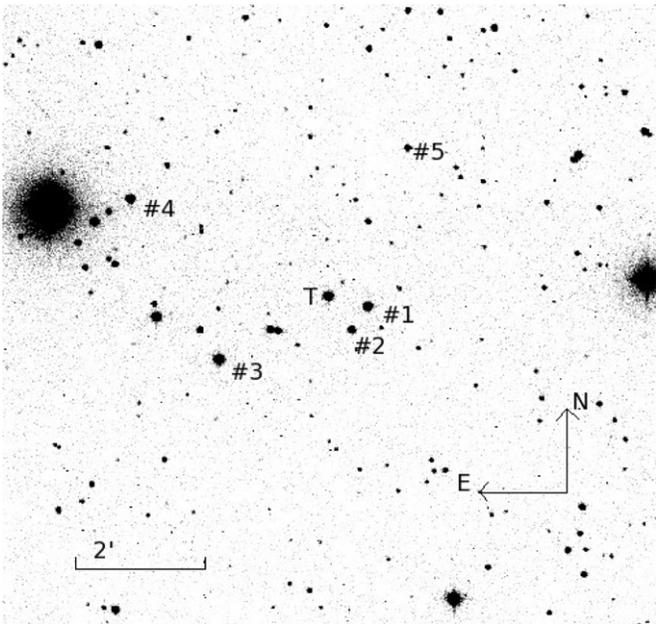


Figure 3. B -band TNT image of MXB 0656–072 taken on 2008 October 08. MXB 0656–072 is marked with “T” and five reference stars are marked with #1 to #5.

performed on MXB 0656–072 using UBV filters, including the observational date, the number of frames and exposure time per filter, and the seeing.

The photometric data reduction was performed using standard routines and aperture photometry packages in IRAF, including bias subtraction and flat-field correction. In order to study the variation of the UBV brightness, we selected five reference stars in the field of view (see Figure 3) to derive the differential magnitude of MXB 0656–072. The results indicate that the reference star #3 is more stable than the other four reference stars. Thus, we selected star #3 as the main reference star and star #4 as check star. The UBV differential magnitudes and errors of MXB 0656–072 in our 2007, 2008, and 2009 observations

are listed in Table 3. The errors are calculated as the square root of the sum of the squares of the measured errors of the target and the main reference star. The typical error for UBV differential magnitudes is within 1%, except for the data on 2009 October 27, when the larger errors may have been due to the bad seeing.

2.3. X-Ray Observations

The All Sky Monitor (ASM, 1.5–12 keV) on board *RXTE* (Levine et al. 1996) and the Burst Alert Telescope (BAT, 15–50 keV) on board *Swift* (Barthelmy 2000) monitored the X-ray activity of MXB 0656–072 since 1998 July and 2005 February, respectively. The X-ray light curves of both *RXTE*/ASM and *Swift*/BAT cover a series of X-ray outbursts since 2007 November (see Figure 4). The positions of the vertical bars at the bottom of the ASM light curve correspond to the beginning time of the *RXTE*-pointed observations.

RXTE public data during the outbursts of MXB 0656–072 between 2007 November and 2008 November were retrieved from the data archive of the High Energy Astrophysics Science Archive Research Center (HEASARC). The *RXTE*/PCA standard2 data with a time resolution of 16 s were used to extract the light curve and the spectrum of the source in the energy range of 3–22 keV. Only Proportional Counter Unit 2 was used for the analysis, because it was always operational during all the pointed observations of MXB 0656–072. We extract Proportional Counter Array (PCA) light curves and spectra when the source has an offset angle lower than $0^{\circ}.02$ and when the earth limb is situated more than 10° away of the source direction. The latest version of PCA background models, which was available at the HEASARC Web site, was used to generate the background light curves and spectra.

The standard mode data of HEXTE cluster B are also used in the analysis. HEASoft script `hxtback` is used to separate the background and source files from the raw FITS files. HEXTE spectra are extracted using `saxextract`, similar to the PCA standard2 data reduction, from the detectors of SpecDet0, SpecDet1, and SpecDet3 in cluster B. Dead time of each

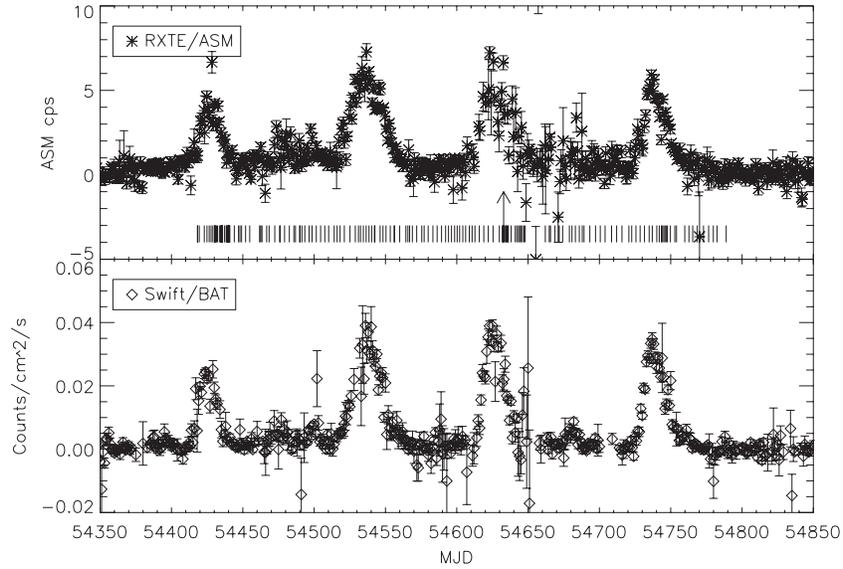


Figure 4. *RXTE/ASM* and *Swift/BAT* light curves displaying several outbursts. The vertical bars in the top panel correspond to the dates of the *RXTE*-pointed observations. The arrow indicates the date of the ObsID 93032-30-01-03.

Table 3
UBV Differential Magnitudes of MXB 0656–072 and the Check Star #4 Using the Main Reference Star #3

MJD	ΔU	ΔU (#4)	ΔB	ΔB (#4)	ΔV	ΔV (#4)
54403.8	0.727 ± 0.006	0.942 ± 0.007	0.887 ± 0.002	0.911 ± 0.002	0.449 ± 0.001	0.960 ± 0.002
54403.8	0.734 ± 0.006	0.944 ± 0.007	0.893 ± 0.002	0.914 ± 0.002	0.455 ± 0.001	0.959 ± 0.002
54404.9	0.728 ± 0.006	0.938 ± 0.007	0.887 ± 0.003	0.906 ± 0.003	0.449 ± 0.003	0.958 ± 0.004
54405.8	0.735 ± 0.005	0.942 ± 0.006	0.886 ± 0.002	0.914 ± 0.002	0.444 ± 0.002	0.958 ± 0.002
54405.8	0.727 ± 0.005	0.934 ± 0.006	0.882 ± 0.002	0.911 ± 0.002	0.442 ± 0.002	0.957 ± 0.002
54744.9	0.520 ± 0.003	0.983 ± 0.004	0.710 ± 0.001	0.911 ± 0.002	0.213 ± 0.001	0.959 ± 0.002
54744.9	0.530 ± 0.003	0.976 ± 0.004	0.709 ± 0.001	0.912 ± 0.002	0.214 ± 0.001	0.959 ± 0.002
54744.9	0.530 ± 0.003	0.980 ± 0.004	0.718 ± 0.002	0.916 ± 0.002	0.219 ± 0.001	0.963 ± 0.002
54745.8	0.545 ± 0.004	0.976 ± 0.004	0.723 ± 0.002	0.913 ± 0.002	0.230 ± 0.001	0.955 ± 0.002
54745.8	0.540 ± 0.004	0.968 ± 0.004	0.731 ± 0.002	0.915 ± 0.002	0.231 ± 0.001	0.954 ± 0.002
54745.8	0.546 ± 0.004	0.967 ± 0.004	0.718 ± 0.002	0.914 ± 0.002	0.233 ± 0.001	0.960 ± 0.002
54747.8	0.532 ± 0.004	0.976 ± 0.004	0.715 ± 0.002	0.919 ± 0.002	0.223 ± 0.001	0.948 ± 0.002
54747.8	0.540 ± 0.004	0.976 ± 0.004	0.728 ± 0.003	0.917 ± 0.004	0.235 ± 0.001	0.958 ± 0.002
54747.8	0.538 ± 0.003	0.973 ± 0.004	0.716 ± 0.002	0.910 ± 0.002	0.232 ± 0.001	0.958 ± 0.002
54747.8	0.532 ± 0.003	0.966 ± 0.004	0.722 ± 0.002	0.919 ± 0.002	0.228 ± 0.001	0.960 ± 0.002
55131.8	0.568 ± 0.035	0.900 ± 0.042	0.784 ± 0.008	0.929 ± 0.008	0.291 ± 0.009	0.986 ± 0.013
55131.8	0.579 ± 0.017	1.039 ± 0.023	0.777 ± 0.012	0.921 ± 0.013	0.315 ± 0.006	0.980 ± 0.008
55131.9	0.598 ± 0.011	0.964 ± 0.013	0.754 ± 0.007	0.914 ± 0.007	0.281 ± 0.004	0.957 ± 0.007
55131.9	0.592 ± 0.010	0.981 ± 0.013	0.763 ± 0.008	0.911 ± 0.009	0.298 ± 0.005	0.968 ± 0.007
55132.9	0.581 ± 0.005	0.985 ± 0.006	0.767 ± 0.004	0.919 ± 0.004	0.291 ± 0.003	0.967 ± 0.004
55132.9	0.578 ± 0.005	0.984 ± 0.006	0.773 ± 0.004	0.924 ± 0.004	0.291 ± 0.003	0.965 ± 0.004

spectra is corrected using `hxtdead`. The background and source spectra are then combined using `sumpha`. The response files are generated and added into the spectra files using the script `hxtresp`.

3. ANALYSIS

3.1. Periodic Variation in X-ray Light Curves

Four continuous X-ray outbursts separated by ~ 100 days were detected by *RXTE/ASM* and *Swift/BAT* telescopes (see Figure 4). These series of outbursts are similar to the periodic Type I X-ray outbursts at periastron passage in Be/X-ray binaries. We selected both X-ray light curves between MJD 54350–54850 to search for periodic variability using the Lomb–Scargle (LS) periodogram method (Scargle 1982). A clear peak at 101.2 days is visible for both ASM (solid line) and BAT (dashed line) LS periodograms (see Figure 5). Combining this value with the spin period of 160.4 s in the $P_{\text{orb}} - P_{\text{spin}}$ di-

agram (Corbet 1986), the source is indeed located in the region of Be/X-ray binaries (see Figure 6). The parameters of the spin and orbital periods of Be/X-ray binaries in Figure 6 are from Liu et al. (2006) and Raguzova & Popov (2005). The position of MXB 0656–072 in the Corbet diagram is close to that of the Be/X-ray binary A0535 + 262 (diamond in Figure 6), which has an orbital period of $P_{\text{orb}} \sim 111.1$ days and a spin period of $P_{\text{spin}} \sim 103$ s (Finger et al. 1996). The 101.2 day period should be the orbital period of MXB 0656–072 and it might produce observational phenomena similar to those of the Be/X-ray binary A0535 + 262.

Both ASM and BAT light curves in Figure 7 are folded, with a period of $P = 101.2$ days and a zero phase epoch of MJD 54408 that corresponds to the beginning of each outburst.

3.2. Optical Variability

Normalized optical spectra of MXB 0656–072 in red and blue regions are shown in Figures 1 and 2, respectively. $H\alpha$,

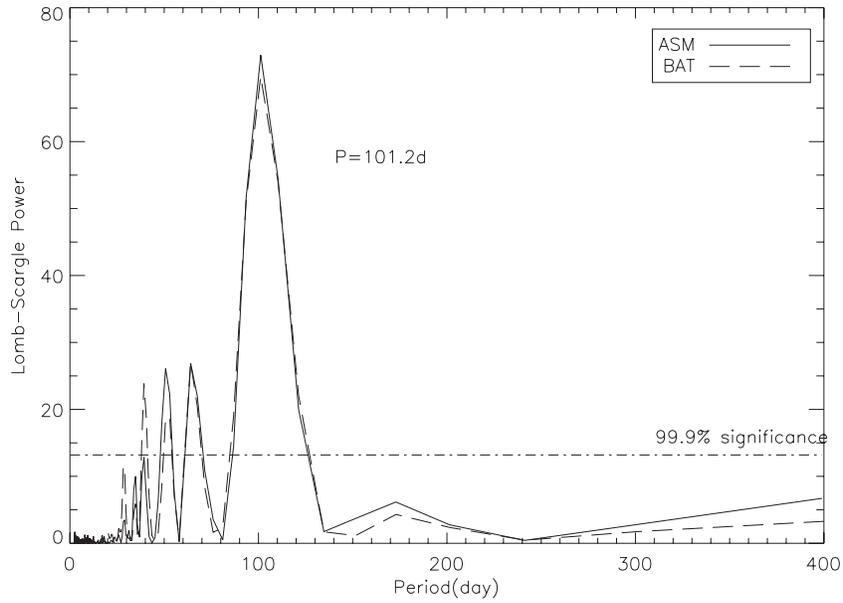


Figure 5. Lomb–Scargle periodogram of ASM (solid) and BAT (dashed) light curves for the source of MXB 0656–072. The horizontal dot-dashed line is the 99.9% confidence level.

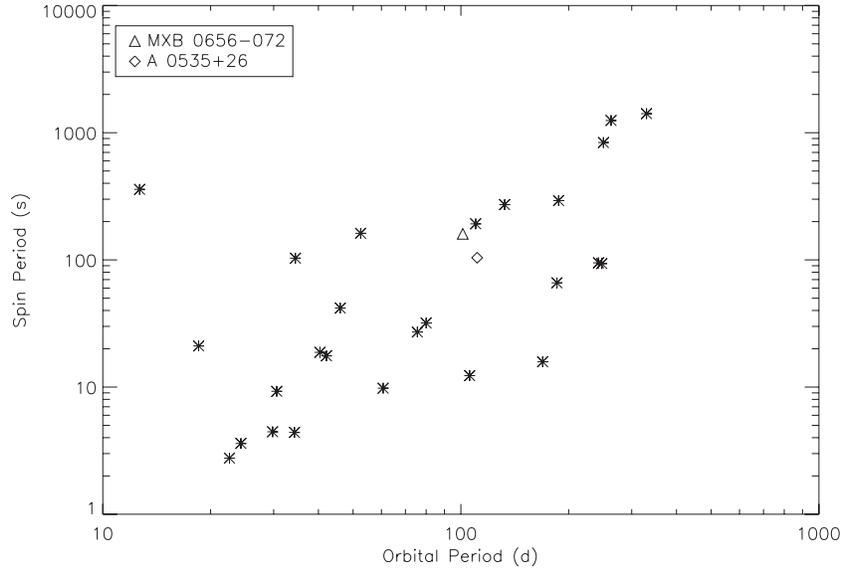


Figure 6. Position of MXB 0656–072 (triangle) in the Corbet diagram. The asterisks are Be/X-ray binaries in catalogs of Liu et al. (2006) and Raguzova & Popov (2005). The diamond corresponds to the position of Be/X-ray binary A0535+262.

He I λ 5875,6678, H β , and H γ are clearly detected. The H α line shows a single-peaked and narrow profile during our observations. All emission lines became very strong just before the X-ray outburst in 2007 November. We plot the EWs(H α), the *UBV* differential magnitudes, and the 1.5–12 keV *RXTE*/ASM light curve in Figure 8. Due to the larger error in the *U*-band magnitude, we do not plot the data of 2009 October 27 (MJD 55131) in Figure 8.

The results show that the strength of the H α line in our 2006 observations became stronger than that of 2005 and it had an extraordinary strength during our 2007 observations, which were taken just before the first X-ray outburst of 2007 November. Our optical spectroscopic observations on 2007 November 16 indicated that the flux of H α kept nearly constant during the X-ray outburst. In comparison with the H α emission lines in our 2005 and 2006 observations, they had almost the same width

of the line wings (see Figure 1(b)), but the H α line during the 2007 observations became much stronger. The H α flux during our 2008 observations dropped to the same level as that in 2006. The optical spectroscopic observations in 2008 were carried out just during the last X-ray outburst. The strength of the H α emission line during our 2009 observations changed little with respect to the strength in 2008. The H β , H γ , and He I lines also had the same variability; for instance, the He I λ 6678 became very strong during our 2007 observations (with an average EW of ~ -1.34 Å) and nearly lost its emission feature in 2005 and 2006 (see Figure 1(a)). In our 2008–2009 observations, the He I λ 6678 line became faint again, but still stronger than in 2005 and 2006.

In 2007, our simultaneous optical photometry and spectroscopy on MXB 0656–072 showed an interesting behavior (see Figure 8): while the H α emission line strongly increased

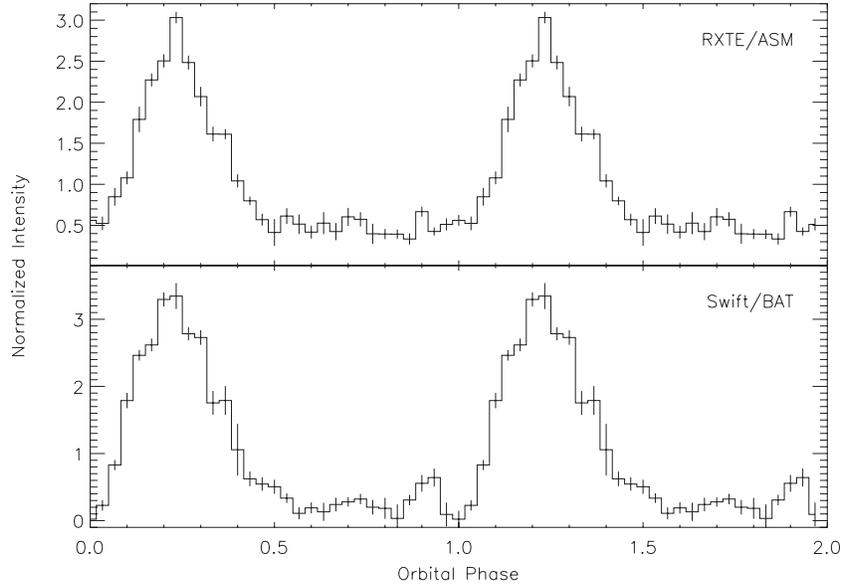


Figure 7. ASM and BAT folded light curves at the period of $P_{\text{orb}} = 101.2$ days. The zero phase epoch (MJD 54408) corresponds to the beginning of the first (2007) outburst.

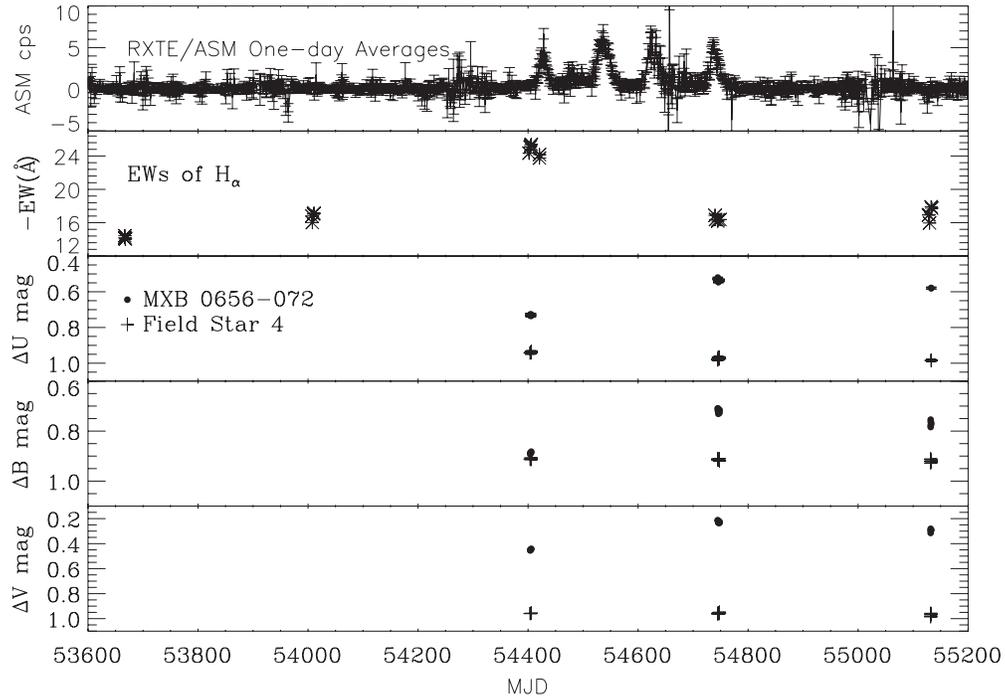


Figure 8. Evolution of EW of $H\alpha$ and UBV differential magnitudes of MXB 0656–072, together with the *RXTE*/*ASM* X-ray light curve in 1.5–12.0 keV. The UBV differential magnitudes of check star 4 are also plotted in the figure. The errors for the EW of $H\alpha$ and UBV differential magnitudes are smaller than the sizes of the corresponding symbols in the figure.

in 2007, the source brightness in UBV decreased by 0.2 mag in 2007 compared to the 2008–2009 observations.

3.3. X-ray Spectral Properties During the X-ray Outbursts

We extract light curves from PCA standard2 data in four different energy bands: 2.3–4.7 keV, 4.7–6.3 keV, 6.3–9.5 keV, and 9.5–21 keV. The soft and hard colors are defined as the count rate ratios 4.7–6.3 keV/2.3–4.7 keV and 9.5–21.0 keV/6.3–9.5 keV, respectively. PCA light curves and the two colors are plotted in Figure 9. The X-ray outbursts are also evident in PCA light curves. The soft and hard colors show the same variability during the X-ray outbursts.

We selected a typical observational ID (ObsID) 93032-30-01-03, which was taken around the time of the peak flux of the third outburst in Figure 4 (the vertical arrow at the bottom of the top panel), to explore the X-ray spectral properties during an X-ray outburst. We modeled the PCA (3–22 keV) and HEXTE (18–80 keV) spectra using ISIS (v1.6.1, Houck & Demicola 2000). Energies lower than 3 keV were ignored due to uncertainties in the background modeling, and a systematic error of 0.5% was added in quadrature to the PCA data in order to take the uncertainties in the response matrix into account. To improve the signal to noise in each channel, we grouped the PCA and the HEXTE spectra to a minimum signal to noise of 5 above a lower bound of 3 keV and 18 keV, respectively.

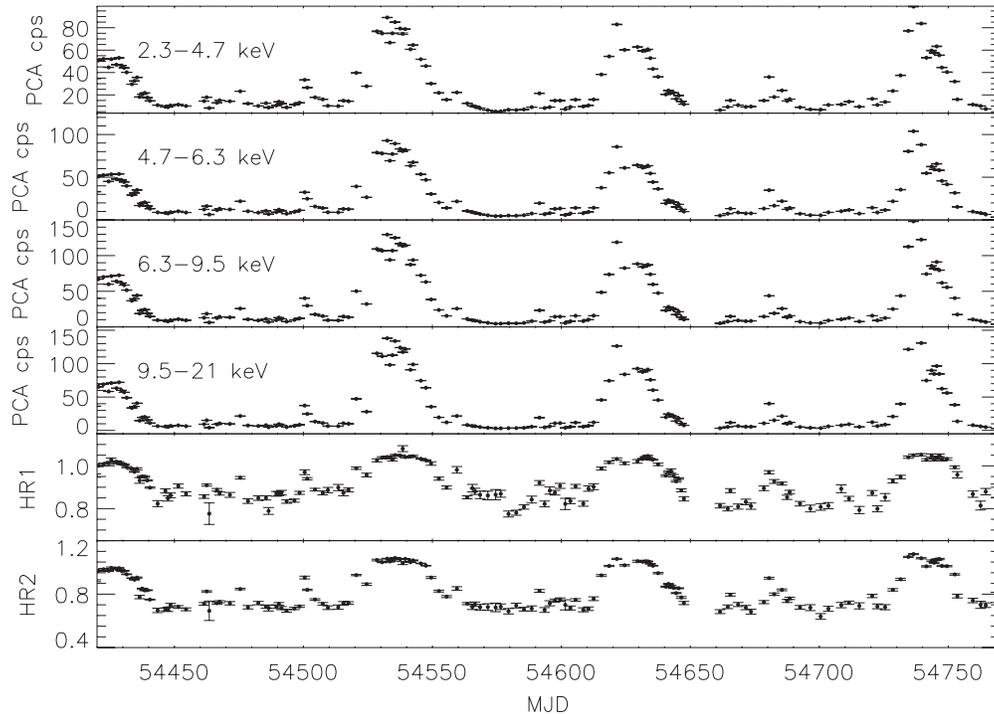


Figure 9. PCA light curves in four different energy bands. The soft and hard colors, defined as $HR1 \equiv 4.7\text{--}6.3\text{ keV}/2.3\text{--}4.7\text{ keV}$ and $HR2 \equiv 9.5\text{--}21.0\text{ keV}/6.3\text{--}9.5\text{ keV}$, respectively, are shown at the bottom. The size of some error bars is smaller than the size of the symbols.

There exists no convincing theoretical model for the shape of the continuum X-ray spectrum in accreting X-ray pulsars. Their X-ray spectrum is generally modeled as a power law with high-energy exponential cutoff. We first tried to fit the PCA-HEXTE spectrum of MXB 0656–072 using a standard ISIS model `cutoffpl`, a power law with a high-energy exponential cutoff,

$$I_{\text{cont}}(E) = \alpha E^{-\Gamma} e^{-E/E_{\text{cut}}}, \quad (1)$$

where α is the photons $\text{keV}^{-1} \text{cm}^{-2} \text{s}^{-1}$ at 1 keV, Γ is the power-law photon index, and E_{cut} is the e -folding energy of exponential cutoff.

Another similar continuum model we tried is a power law with a high-energy rollover (PLCUT; White et al. 1983), which was applied to MXB 0656–072 during the 2003 X-ray outburst by McBride et al. (2006). The analytic form of the PLCUT model is

$$I_{\text{cont}}(E) = \alpha \times E^{-\Gamma} \begin{cases} 1 & E \leq E_{\text{cut}} \\ e^{-(E-E_{\text{cut}})/E_{\text{fold}}} & E > E_{\text{cut}}, \end{cases} \quad (2)$$

where E_{cut} and E_{fold} are the cutoff and folding energies, respectively.

Each fit shows in the residuals a pseudo-P Cygni profile near 10 keV, which has been reported by McBride et al. (2006) and found in other accreting X-ray pulsars (Her X-1, GS 1843+00; Coburn et al. 2002). They concluded that the pattern of the residuals around 10 keV was due to the continuum model. In the PLCUT model, the energy cutoff should be smoothed by including a Gaussian absorption line with an energy-dependent width at the continuum cutoff energy (McBride et al. 2006). Here, we choose to use `cutoffpl` with a low-energy absorption, as it is the simplest model, to fit the X-ray continuum emission for the spectrum of MXB 0656–072. A normalization constant (1 for the PCA data and a free parameter for the HEXTE data)

was added to the model to get a good agreement between these two instruments. A bad reduced Chi-square $\chi^2_{\nu} = 21.5$ for 53 degrees of freedom (dof) was obtained and the residuals are shown in panel (b) of Figure 10. Since the feature around 10 keV might not be real, we will not discuss its physical nature in this paper. In order to remove the systematic feature in the residuals at ~ 10 keV, a Gaussian absorption line (`gabs`) function,

$$\text{GABS}(E) = \exp\left(-\frac{\tau}{\sqrt{2\pi}\sigma}\right) \times \exp\left(-\frac{(E - E_l)^2}{2\sigma^2}\right), \quad (3)$$

was used in the model and improved the fit from a reduced Chi-square of 21.5 for 53 dof to 8.9 for 50 dof, where τ is the optical depth, σ is the line width, and E_l is the line energy. The other obvious feature in the panel (b) of Figure 10 should be an iron line at ~ 6.4 keV, which was described by the Gaussian-shaped model `gaussian` in ISIS. Due to the lower energy resolution of PCA (less than 18% at 6 keV), the width of the iron line in the model was fixed at 0 keV. After this feature was added in the model, a better reduced Chi-square of 1.23 for 48 dof was obtained. Thus, the F -statistic is 156.7 and the F -test probability is about 9.1×10^{-22} , which indicates that it is reasonable to add a Gaussian component in the model. Protassov et al. (2002) suggested that the F -test statistic cannot be used to test for the presence of a line and they proposed a Monte Carlo simulation approach to calculate the approximate statistical significance. We followed this method to calibrate the sample distribution of the F -statistic for the spectrum of MXB 0656–072. We simulated 10,000 spectra with ISIS from the null model (a cutoff power law with the low-energy absorption and a Gaussian absorption feature around 10 keV) at its best-fit parameters. Each faked data were grouped exactly as we did for the real data and fitted with the alternative model (which included the Gaussian line component). The F -statistic between the null and the alternative model was calculated and their distribution was shown in Figure 11. The maximum value of the F -statistic from

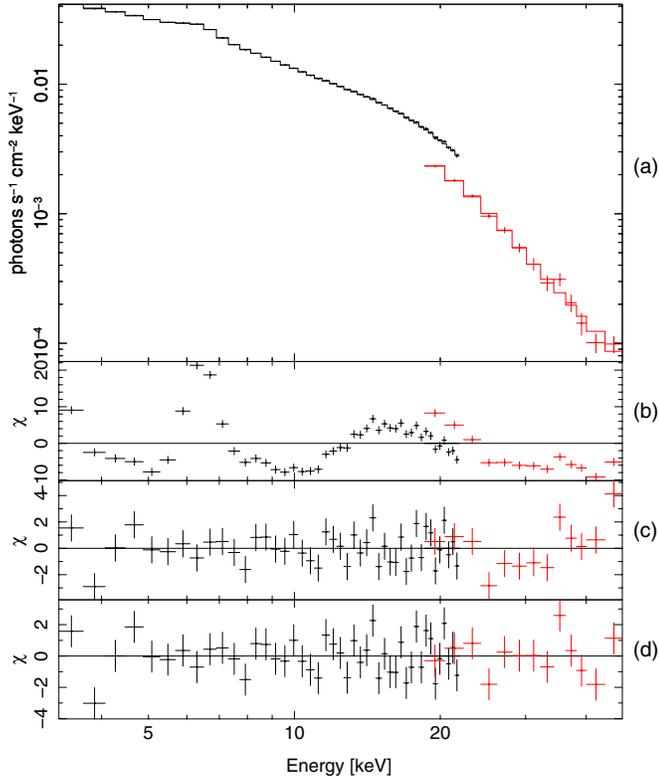


Figure 10. Best-fit PCA and HEXTE spectra for the ObsID 93032-30-01-03 modeled with a cutoff power law with low-energy absorption, a Gaussian line around 6.4 keV, and a cyclotron resonance scattering feature around 30 keV. The panels (b), (c), and (d) are the residuals only for the `cutoffpl` with the low-energy absorption, the residuals without `cyc1abs`, and the best-fit residuals, respectively. The residuals of the fit are in terms of σ values.

(A color version of this figure is available in the online journal.)

the 10,000 faked spectra is much less than the F -statistic value of 156.7 calculated from the real data. Therefore, we concluded that the ion line around 6.4 keV was detected at the greater than 99.99% confidence level in the spectrum of MXB 0656–072.

We found a broad absorption feature around 30 keV in the residuals of the last fitting which showed in the panel (c) of Figure 10. This feature should be the cyclotron resonance scattering feature around ~ 30 –40 keV reported by McBride et al. (2006) and it was modeled with the `cyc1abs` function in ISIS,

$$\text{CYCLABS}(E) = \exp\left(-D_{\text{cycl}} \frac{(W_{\text{cycl}} E / E_{\text{cycl}})^2}{(E - E_{\text{cycl}})^2 + W_{\text{cycl}}^2}\right), \quad (4)$$

where, E_{cycl} , D_{cycl} , and W_{cycl} are the cyclotron energy, depth, and width, respectively (Mihara et al. 1990; Makishima et al. 1990). Here, the cyclotron width W_{cycl} was fixed at 10 keV in the model. A good reduced Chi-square of 1.01 for 46 dof was obtained and the residuals were shown in panel (d) of Figure 10. With the same method as the iron line, the detection significance of the cyclotron resonant absorption feature was also at greater than 99.99% confidence level. The best-fit spectral parameters for the ObsID 93032-30-01-03 are listed in Table 4, together with the MJD, the observational date, the exposure time, and the model flux in 3–20 keV at a distance of 3.9 kpc (McBride et al. 2006). The errors were estimated at the 90% confidence level ranges of the parameters.

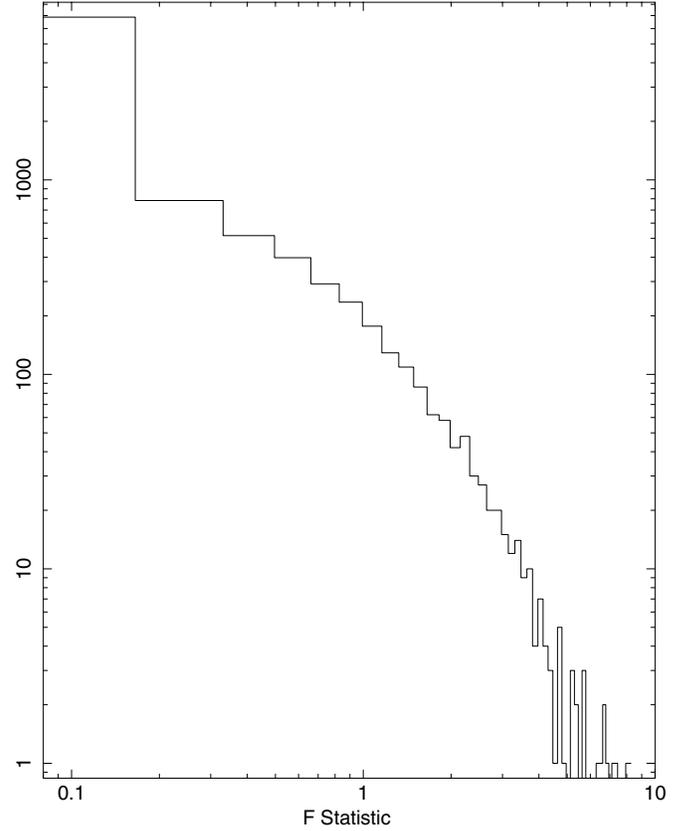


Figure 11. F -statistic distribution between the null model (see the text) and an alternative model with the addition of a Gaussian line around 6.4 keV for 10,000 simulated spectra.

Table 4
Best-fit Parameters for PCA-HEXTE Spectra
of MXB 0656–072

Parameter	Value
ObsID	93032-30-01-03
MJD	54632.88
Date	2008 Jun 15
Exposure (s)	4912
N_{H} (10^{22} cm^{-2})	$1.50^{+0.53}_{-0.50}$
E_{Fe} (keV)	$6.44^{+0.05}_{-0.03}$
FeNorm^a	$4.3^{+0.4}_{-0.4}$
E_{GABS} (keV)	$10.45^{+0.22}_{-0.25}$
σ_{GABS} (keV)	$2.83^{+0.64}_{-0.53}$
τ_{GABS}	$1.03^{+0.77}_{-0.41}$
Γ	$0.37^{+0.18}_{-0.19}$
E_{fold} (keV)	$9.45^{+2.02}_{-1.53}$
α^b	$0.10^{+0.02}_{-0.02}$
E_{cycl} (keV)	$30.2^{+1.9}_{-2.9}$
D_{cycl}	$0.5^{+0.3}_{-0.2}$
Factor ^c	$0.59^{+0.01}_{-0.01}$
χ_{red}^2 (dof)	1.01(46)
Luminosity (3–22 keV) ^d	6.6

Notes. All errors represent 90% confidence intervals.

^a $10^{-3} \text{ photons cm}^{-2} \text{ s}^{-1}$.

^b $\text{photons keV}^{-1} \text{ cm}^{-2} \text{ s}^{-1}$ at 1 keV.

^c The normalization constant for the HEXTE data.

^d In units of $10^{36} \text{ erg s}^{-1}$ for a distance of 3.9 kpc (McBride et al. 2006).

4. DISCUSSION

We have analyzed the optical and X-ray observations of MXB 0656–072 during different X-ray states. A 101.2 day period has been inferred from the *RXTE*/ASM and *Swift*/BAT light curves thanks to a series of X-ray outbursts which occurred between MJD 54350 and MJD 54850. Such a 101.2 day period was very likely the orbital period of MXB 0656–072. The position of this 101.2 day period and its 160.4 s spin period in the Corbet diagram support the Be/X-ray binary nature of the system. MXB 0656–072 has similar orbital parameters as the Be/X-ray binary A 0535+26, which has a 111 day orbital period and a 103 s spin period, and they almost occupy the same position in the Corbet diagram (see Figure 6). The folded *RXTE*/ASM and *Swift*/BAT light curves of MXB 0656–072, with an orbital period of $P_{\text{orb}} = 101.2$ days, show a complex outburst profile (see Figure 7). The duration of the outburst covers nearly half of one orbital phase (~ 50 days). Using the third Kepler law, we can estimate the semimajor axis of the binary system with $a^3 = P_{\text{orb}}^2 \times G(M_* + M_X)/4\pi^2$, where M_* and M_X are the masses of the optical companion and neutron star, respectively. Since MXB 0656–072 (O9.7 Ve) and A0535+26 (O9.7 IIIe) have similar optical counterparts, we adopt $M_* = 20 M_\odot$ and $R_* = 15 R_\odot$ (Grundstrom et al. 2007). The mass of the compact object is fixed to $M_X = 1.4 M_\odot$, typical of a neutron star. These values yield $a \sim 253 R_\odot = 16.9 R_*$. The H α emission line is generally found to arise in a region about $10R_*$ from the central Be star (Slettebak et al. 1992). In order to capture material and cause an X-ray outburst, the neutron star should pass within this region at the periastron point. Hence, the separation between the neutron star and the Oe star at the periastron point should be smaller than the H α emission region, given by $a(1 - e) \sim 10R_*$. Thus, we can estimate that the orbital eccentricity of MXB 0656–072 should be $e \sim 0.4$, which is comparable to the eccentricity of A0535+26, i.e., $e = 0.47$ (Finger et al. 1996).

We have followed MXB 0656–072 spectroscopically since 2005 and also photometrically since 2007 at optical wavelengths at Xionglong Station, NAOC. The strong H α emission line showed a maximum EW of $\sim 25 \text{ \AA}$ during our 2007 observations, which were taken just before the X-ray outburst in 2007 November. A strong He I $\lambda 6678$ line was also observed during the 2007 observations. The photometric observations showed that the *UBV* magnitudes during our 2007 observations were about 0.2 mag larger with respect to those in our following observations (see Table 3 and Figure 8). Spectroscopy of 2007 November 16 (during the rising phase of the first X-ray outburst) indicated that the H α still had a very strong emission during the X-ray outburst. It is generally believed that the H α emission line in Be stars is formed in the entire circumstellar disk (Slettebak et al. 1992), while only the innermost part of the disk contributes significantly to the continuum flux. Due to the higher ionization potential energy, the formation region of the He I $\lambda 6678$ line should be smaller than the nearby continuum region (Stee et al. 1998). The consistency among the variability of the brightness in the *UBV* bands in MXB 0656–072 suggests that the changes of the optical continuum emission might be caused by the physical changes in the circumstellar disk around the Oe star (Janot-Pacheco et al. 1987). The strong H α emission corresponds to a more extended circumstellar disk, while the decrease of the optical brightness indicates the dilution of the inner disk. Moreover, the increase of the He I $\lambda 6678$ emission in the 2007 observations should be connected with the formation of a denser inner disk.

To explain all the observational phenomena, a unified physical model should be assumed in MXB 0656–072. Observational results (Rivinius et al. 2001) and theoretical calculations (Meilland et al. 2006) suggest that after an outburst, a low-density region seems to develop around the Be star. Rivinius et al. (2001) and Meilland et al. (2006) suggested that the outburst might be connected with the increased mass loss or mass ejection from the Be star. Some weeks to months after the outburst, the stellar radiation pressure gradually excavates the inner part of the disk and a low-density region could develop around the Be star and slowly grow outward (Rivinius et al. 2001). With the vacuation of the inner disk, the optical continuum emission decreases and an increase in *UBV* magnitudes will be observed. After the outburst, material is transferred into the disk and a more extended circumstellar disk should be formed, which produces the stronger H α emission from the system. Therefore, we suggest that a mass ejection event had taken place in MXB 0656–072 before our 2007 observations and a low-density region was developing during that period. The He I $\lambda 6678$ became stronger when the optical continuum emission was decaying, which indicated that a larger quantity of material close to the stellar surface should be present when a low-density region was developing in the inner part of the disk. The new disk material might be ejected from the star by the subsequent mass ejection or reaccreted into the inner region after the supply of material from the star to the disk has been turned off (Clark et al. 2001).

With the expansion of the circumstellar disk after the outburst, the outer part of the disk interacts with the neutron star around its periastron passage. A part of the material in the disk is accreted onto the neutron star and an X-ray outburst is triggered. This could explain the first X-ray outburst in 2007, which occurred in a decline phase of the optical brightness. A considerable fraction of the angular momentum would be transferred onto the neutron star in the system, which may account for the simultaneous spin-up detected by *Fermi*/Gamma-ray Burst Monitor⁵ during the last X-ray outburst in Figure 4. A similar phenomenon was also observed in other Be/X-ray binaries, such as 4U 1145-619 (Stevens et al. 1997), A0535+26 (Clark et al. 1999; Yan et al. 2012), and 4U 0115+63 (Reig et al. 2007). All these X-ray outbursts seem to take place in the decline phase of the optical emission.

During the first X-ray outburst in 2007, the H α emission kept nearly unchanged. The motion of the neutron star could not influence the size of the H α emission region instantaneously. With the extension of the ejected material, the density in the outer region of the disk would increase and a much stronger peak flux was observed in the second and third X-ray outbursts (see Figure 4). The last X-ray outburst in 2008 occurred between MJD 54720 and MJD 54760. Our 2008 optical observations, taken during the last X-ray outburst, show that the H α emission level and the *UBV* brightness returned to their levels of 2006. There were no more changes in H α emission and the *UBV* brightness between our 2008 and 2009 observations. However, no X-ray outburst has been observed since the last one in 2008, indicating that the material ejected in 2007 has been dispersed into the outer region of the system. The circumstellar disk in the equatorial region of Oe star became denser and larger at the early stage of the mass ejection in the end of 2007 and a stronger H α emission was observed during our 2007 observations. Due to the gravitational effect of the neutron star, the size of the circumstellar disk could not increase continuously and it should

⁵ <http://gammaray.msfc.nasa.gov/gbm/science/occultation/>

have been truncated (Okazaki & Negueruela 2001). The size of the truncated disk after the last X-ray outburst in 2008 should be smaller than the distance to the first Lagrangian (L1) point at periastron and the H α emission was in a stable level after our 2008 observations.

Another large, extended Type II X-ray outburst was observed with *RXTE* in MXB 0656–072 from 2003 October to 2004 January (McBride et al. 2006). The duration of that Type II X-ray outburst is longer than the X-ray outbursts observed between 2007 November and 2008 November. It lasted about 100 days between MJD 52924 and MJD 53026. There were no systematic optical observations during 2003 X-ray outburst and we could not know what happened in MXB 0656–072. According to its long X-ray outburst duration (nearly the whole orbital period), we speculate that a much stronger mass ejection might have taken place before its 2003 outburst. Due to the spectacular changes of the stellar wind environment during the Type II X-ray outburst, the disk size would increase very rapidly and the outer disk material could interact with the neutron star before or after it approaches to the periastron point. Therefore, the Type II X-ray outburst could occur at any orbital phase and last a longer time than the Type I X-ray outburst.

The PCA + HEXTE spectra of MXB 0656–072 during the 2007–2008 X-ray outbursts could be well fitted by a cutoff power-law model with a low-energy absorption, together with an iron fluorescent line at ~ 6.4 keV and a cyclotron feature at ~ 30.2 keV. Unlike the model suggested by McBride et al. (2006) for the X-ray spectrum during the 2003 X-ray outburst, a blackbody component is not necessary in our model for the X-ray spectrum during the 2007–2008 X-ray outbursts. The detection significance for the iron line around 6.4 keV and the cyclotron resonance scattering feature around 30 keV is at a larger than 99.99% confidence level. The X-ray luminosity in the 3–22 keV could be calculated as 6.6×10^{36} erg s $^{-1}$, assuming a distance of 3.9 kpc (McBride et al. 2006). Once the energy of the cyclotron resonance line was known, the strength of the magnetic field in the surface of the neutron star could be estimated as

$$B = \frac{E_{\text{cycl}}}{11.6 \text{ keV}} \times (1+z) \times 10^{12} \text{ G}, \quad (5)$$

where E_{cycl} is the energy of the fundamental cyclotron resonant scattering line and z is the gravitational redshift, respectively (Coburn et al. 2002). Assuming $E_{\text{cycl}} = 30.2^{+1.9}_{-2.9}$ in Table 4 is the fundamental cyclotron line and $z = 0.3$ for a typical neutron star with mass of $1.4 M_{\odot}$ and radius of 10 km (McBride et al. 2006), the magnetic field could be calculated by Equation (5) to be $3.4^{+0.2}_{-0.3} \times 10^{12}$ G. The same variability of the soft and hard colors during the X-ray outbursts suggests no overall changes in the spectral shape of MXB 0656–072 in the 2.3–21.0 keV band. The X-ray flux during the outbursts should only be connected with the changes of the mass accretion rate onto the neutron star around the periastron point.

5. CONCLUSIONS

We presented and analyzed the optical and X-ray observations of MXB 0656–072 during different X-ray outbursts. A 101.2 day orbital period was found from a series of X-ray outbursts of the system between MJD 54350 and MJD 54850. The position of this orbital period and the 160.4 s spin period in the Corbet diagram confirms the Be/X-ray binary nature of MXB 0656–072. The anti-correlation between the H α emission

and the *UBV* brightness during our 2007 observations indicates that a mass ejection event might have taken place before the first X-ray outburst in 2007 November. A low-density region was developed around the Oe star in MXB 0656–072 after the mass ejection event, which could explain the decrease of the optical brightness during the 2007 observations. Material was transferred into the circumstellar disk during the mass ejection event and a more extended disk formed after the mass outburst, which should be the reason for the strong H α emission before the 2007 X-ray outburst. With the outward motion of the disk, the neutron star interacted with the outer part of the disk around the periastron passage and a series of transient X-ray outbursts could be triggered. According to the long duration of the 2003 Type II X-ray outburst, we speculate that a much stronger mass ejection event should have taken place before the 2003 outburst. The ejected material could fill all the orbital regions of the neutron star.

The PCA-HEXTE spectra during the 2007–2008 X-ray outbursts could be well fitted by a cutoff power law with low-energy absorption. A Gaussian iron line around 6.4 keV and a cyclotron resonance scattering line around 30 keV are also considered in the model. The same variability of the two X-ray colors in the 2.3–21 keV band indicated no overall changes in the spectral shape during the X-ray outbursts, which might only be connected with the changes of the mass accretion rate onto the neutron star.

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